



# A coupled model for the formation of an active region corona

Feng Chen

*Max Planck Institute for Solar System Research  
Georg-August-Universität Göttingen*

*chen@mps.mpg.de*

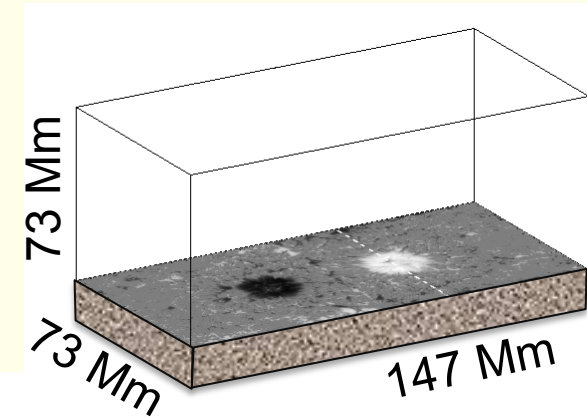
H. Peter, S. Bingert, R. Cameron,  
M. Schüssler, and M. C. M. Cheung

# Coronal model driven by emerging flux simulation

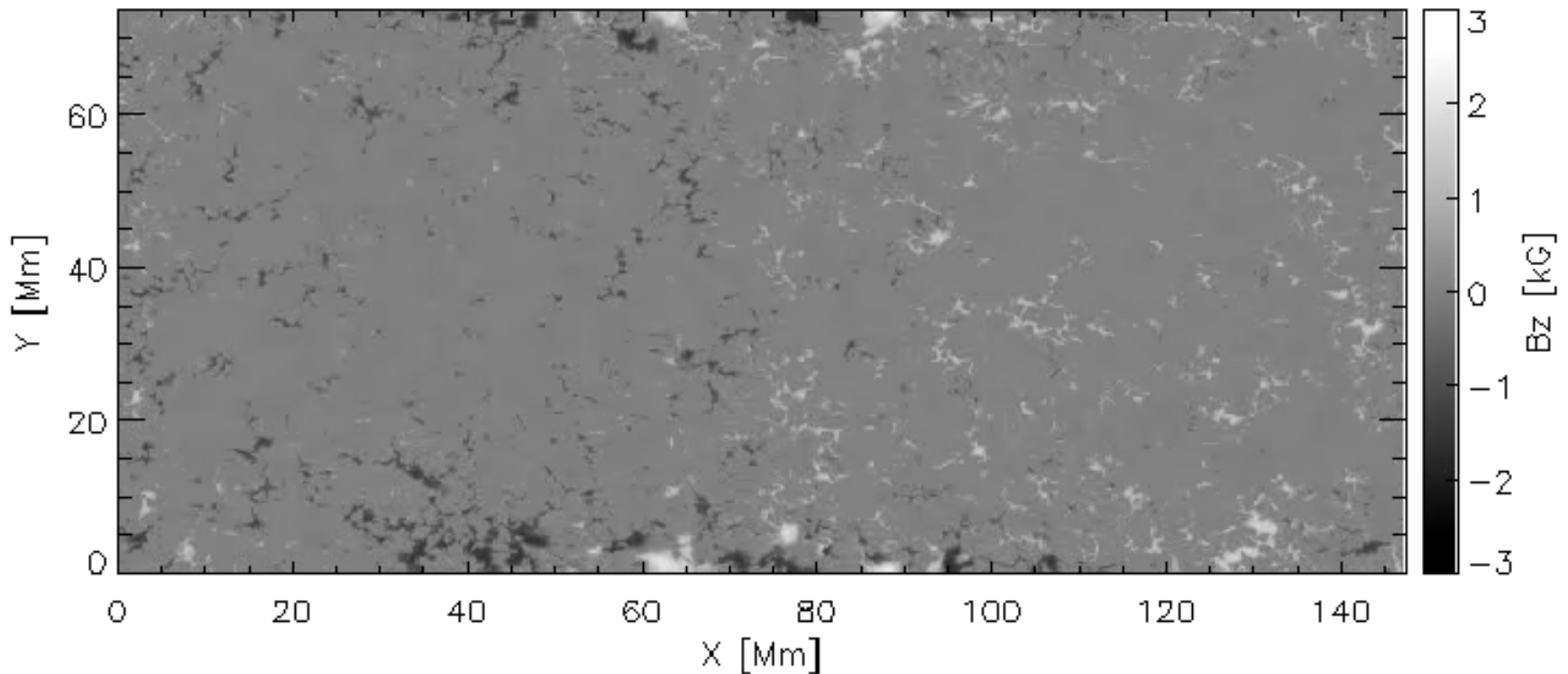
## flux-emergence simulation

from / similar to Cheung et al (2010) ApJ 720, 233

- flux rope rises from bottom and breaks through surface
- pair of sunspots



20h 58.97min



# Coronal model driven by emerging flux simulation

## flux-emergence simulation

from / similar to Cheung et al (2010) ApJ 720, 233

– flux rope rises from bottom  
and breaks through surface

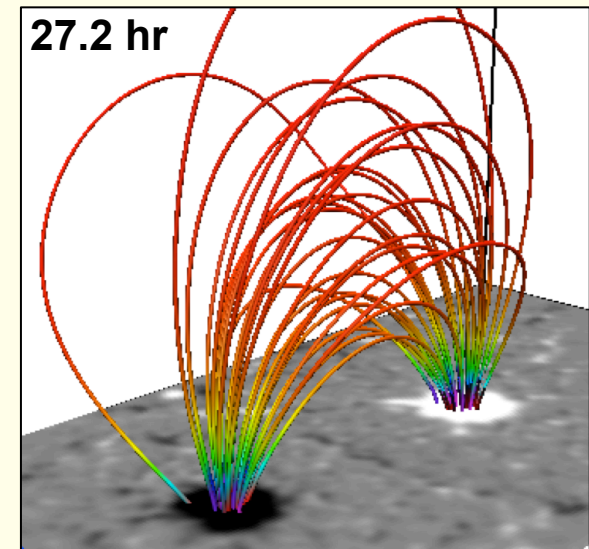
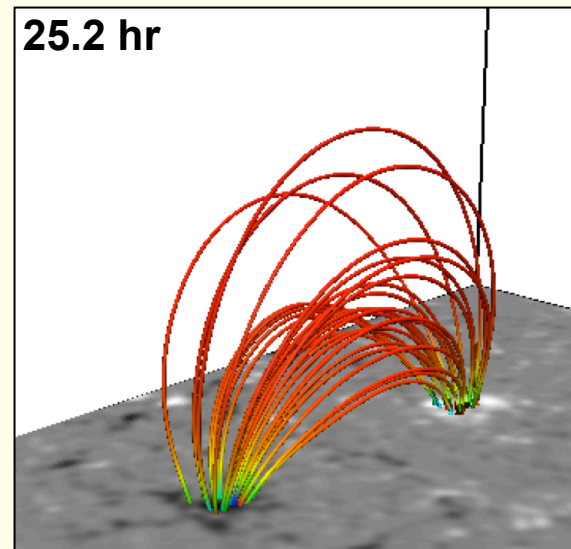
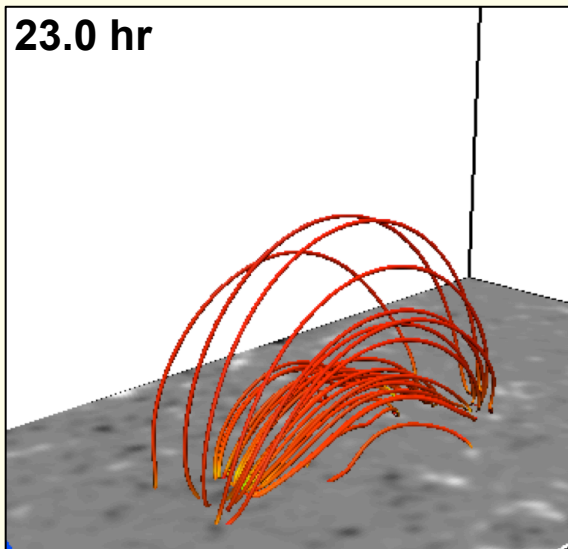
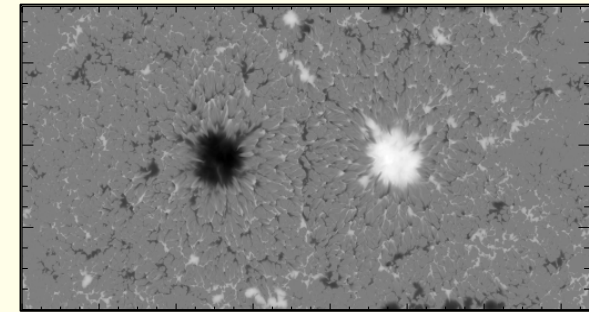
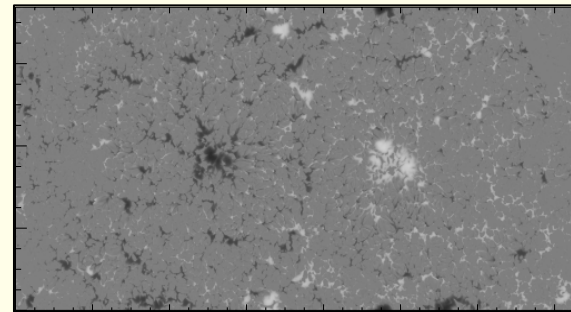
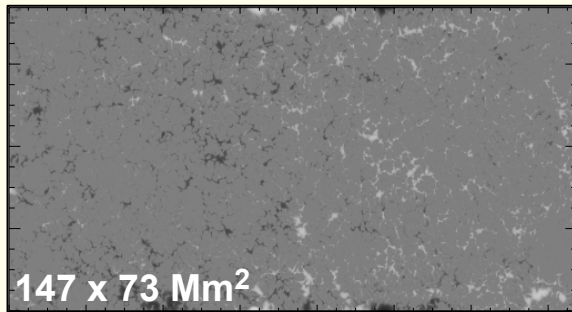
→ pair of sunspots

## coronal simulation

– use photospheric layer  $(T, \rho, v, B)$   
as time-dependent lower boundary

→ magnetic field expands

→ coronal loops form



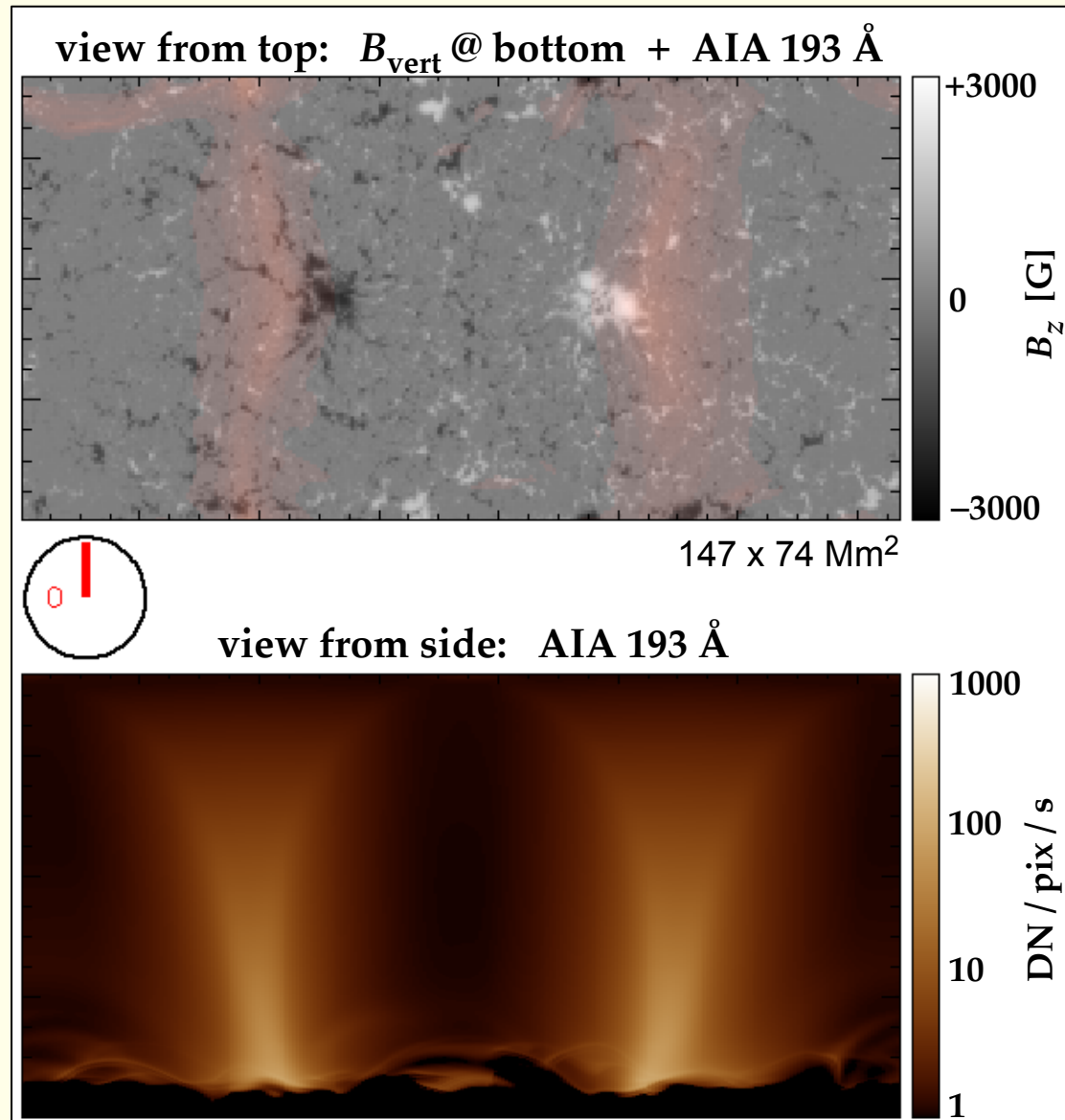
# Coronal model driven by emerging flux simulation

synthesized coronal emission ( $1.5 \cdot 10^6$  K)

- ▶ loops form at different places at different times
- ▶ loops appear quickly ( $< 5$  min)
- ▶ loop footpoints are in sunspot periphery
- ▶ at times appears with constant cross section  
( Peter & Bingert 2012, A&A, 548A, 1)

Question:

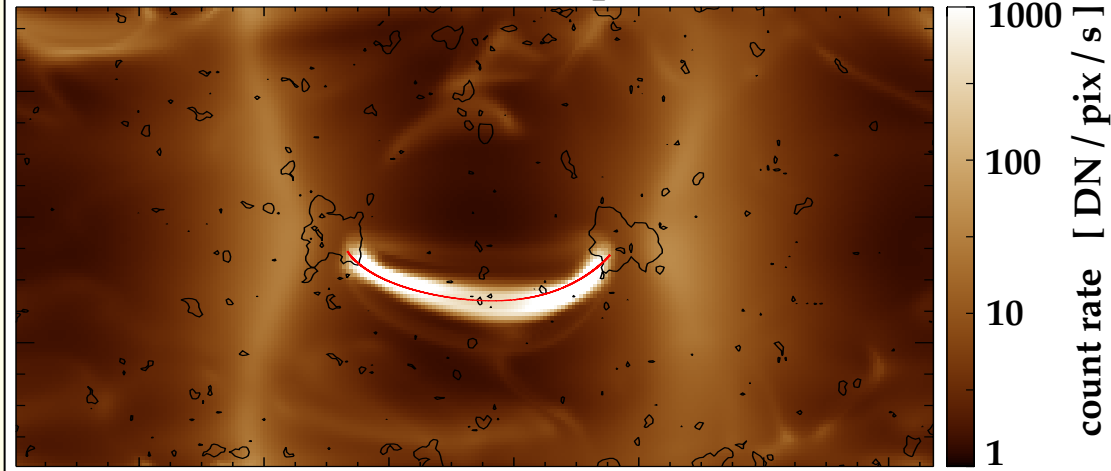
- ▶ Why, where, when does the loop form?



# Coronal loop and energy input

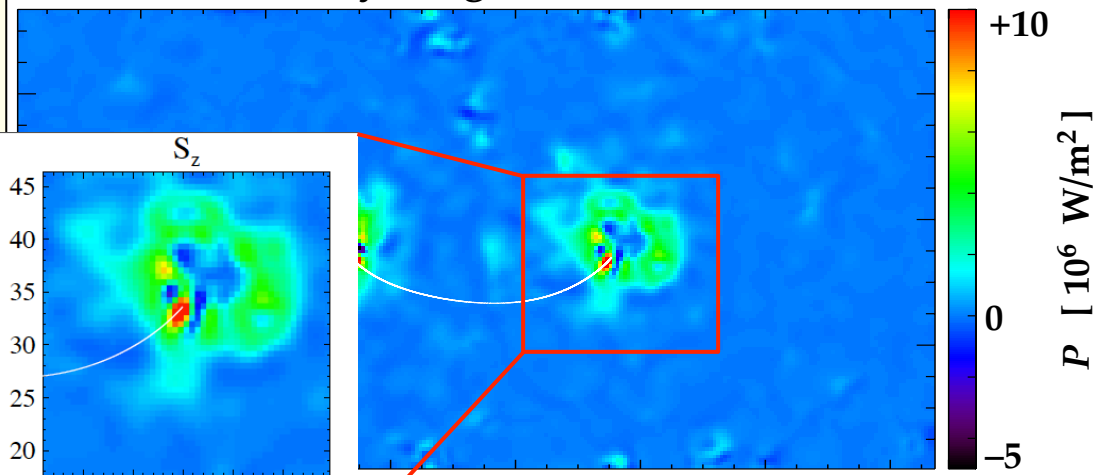
- ▶ loops form at different places at different times
- ▶ loops appear quickly ( $< 5$  min)
- ▶ loop footpoints are in sunspot periphery
- ▶ at times appears with constant cross section ( Peter & Bingert 2012, A&A, 548A, 1)

coronal emission from top: AIA 193 Å

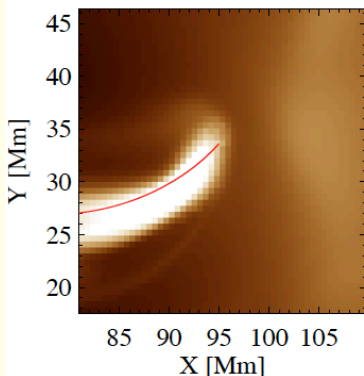


147 x 73 Mm<sup>2</sup>

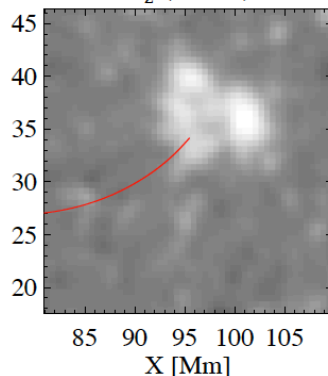
vertical Poynting flux near bottom



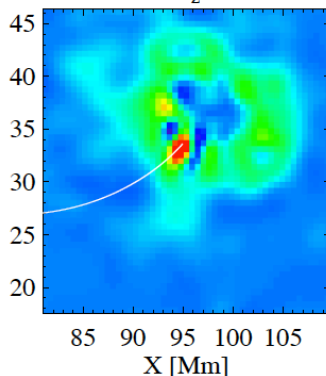
AIA 193



$B_z$  (Z = 0)



$S_z$



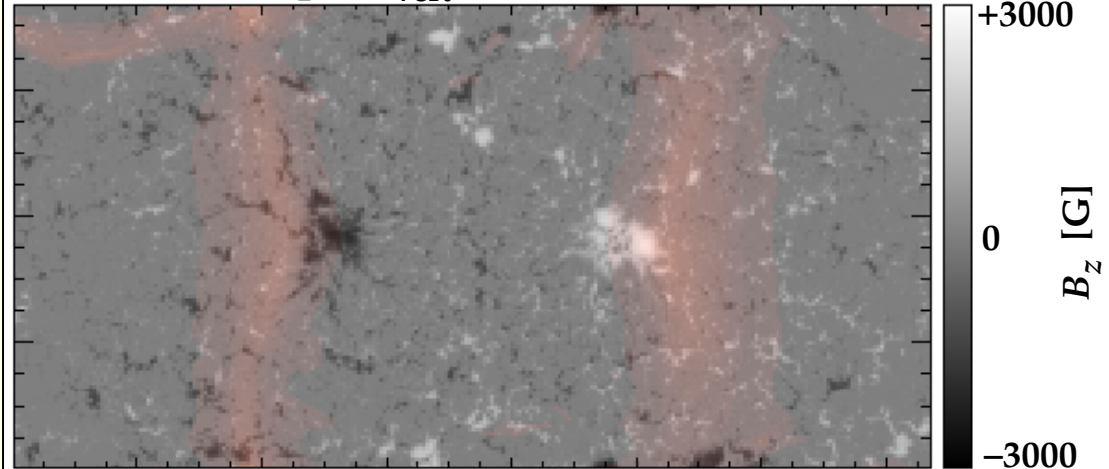
$P$  [ $10^6$  W/m<sup>2</sup>]

# Expansion and fragmentation of loop

- ▶ loops form at different places at different times
- ▶ loops appear quickly ( $< 5$  min)
- ▶ loop footpoints are in sunspot periphery
- ▶ at times appears with constant cross section  
( Peter & Bingert 2012, A&A, 548A, 1)

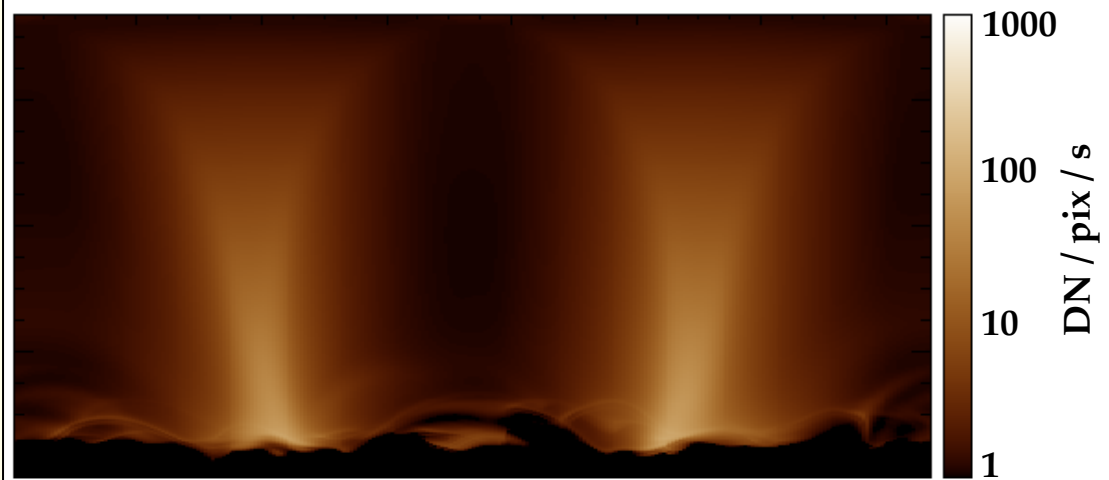
synthesized coronal emission ( $1.5 \cdot 10^6$  K)

view from top:  $B_{\text{vert}}$  @ bottom + AIA 193 Å



$147 \times 74 \text{ Mm}^2$

view from side: AIA 193 Å

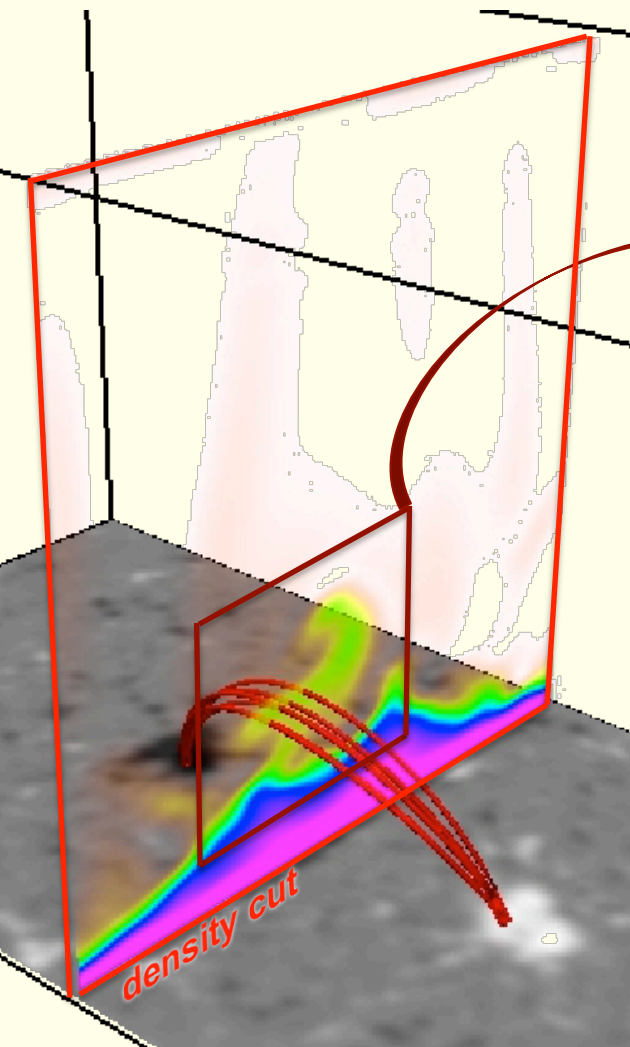


34 min out of 7 hrs

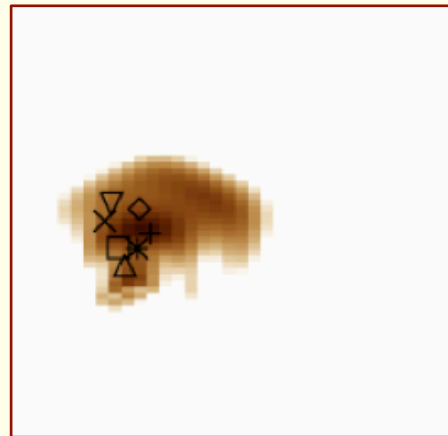
Question:

- ▶ Expansion and fragmentation of loop

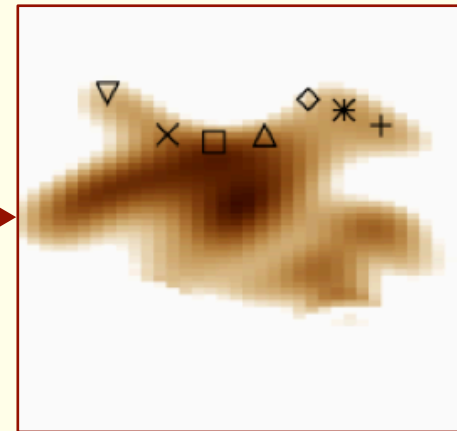
# Defining a magnetic tube



cross section in the loop mid-plane



10 min  
later



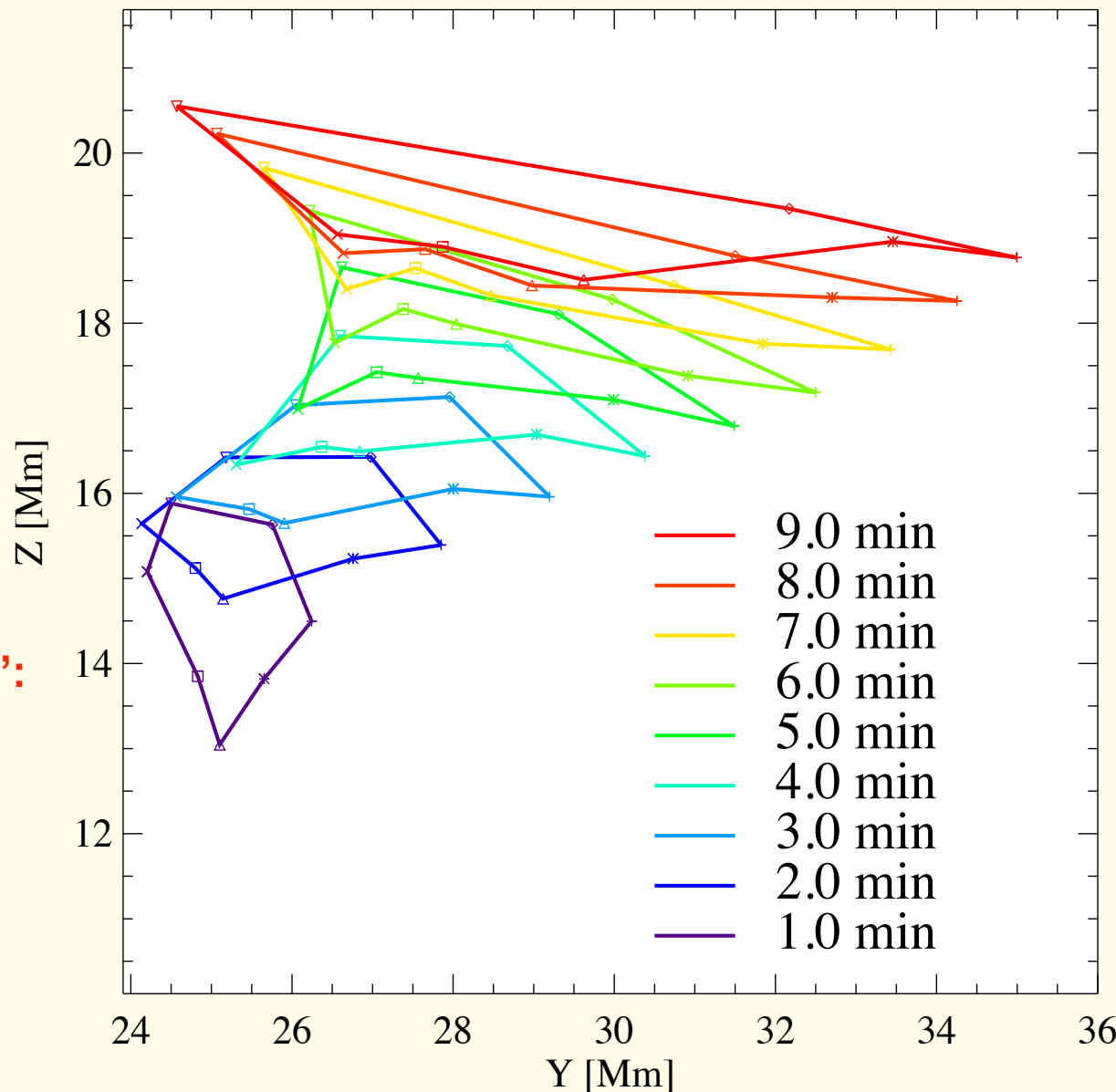
AIA 193 emission

# Expansion of the magnetic tube

► Cross section of the magnetic tube

► Cross section slightly increasing

► **Non-isotropic “expansion”:**  
deformation of  $B$  during  
flux emergence

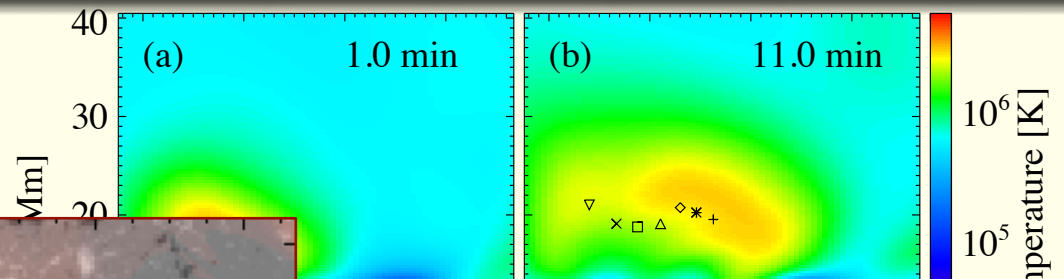




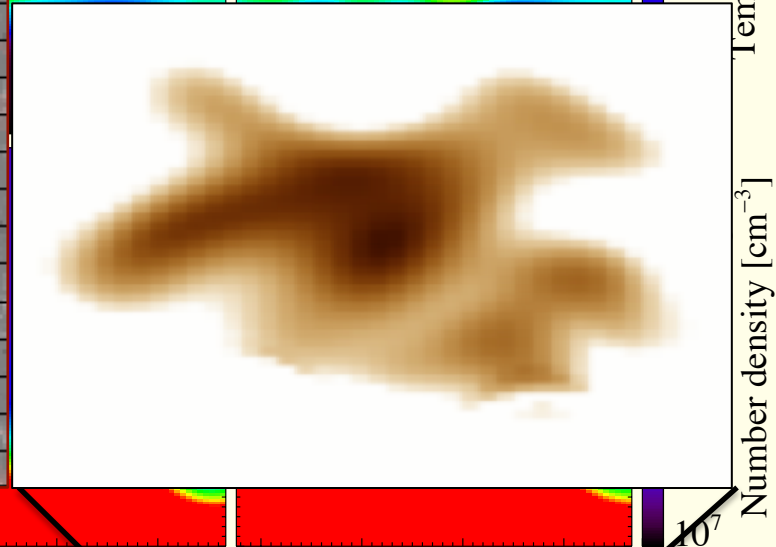
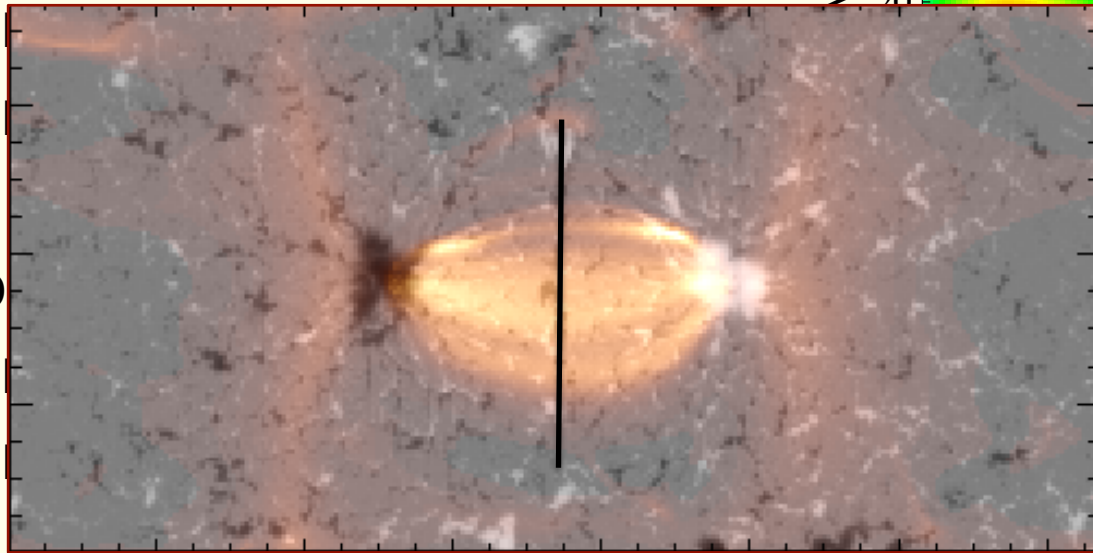
# Loop fragmentation

Temperature

- "Hot shell"



D

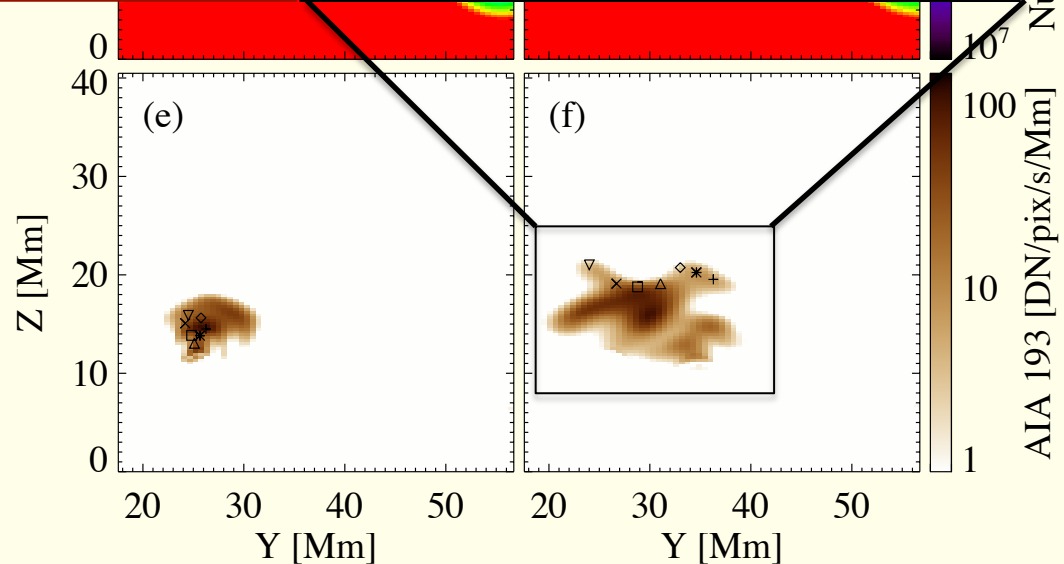


Emissivity

- Convolution of  $T$  and  $\rho$

- Shows finer structure than both  $T$  and  $\rho$

- Fragmentation shows also for l.o.s. integration



# Conclusions

- ▶ Flux emergence through bottom boundary drives evolution of AR corona
- ▶ Heating of corona by dissipation of induced currents
- ▶ Strong horizontal expansion of the magnetic tube
- ▶ Fragmentation of the loop seen in coronal emission
- ▶ loop is truly 3D in nature

